

ξ^1 CMa: An Extremely Slowly Rotating Magnetic B0.7 IV Star

Matt Shultz^{1,2,3}, Gregg Wade³, Thomas Rivinius¹
Wagner Marcolino⁴, Huib Henrichs⁵, Jason Grunhut¹
and the MiMeS Collaboration

¹European Southern Observatory; ²Queen's University, Canada

³Royal Military College, Canada ⁴Observatório do Valongo, UFRJ, Brazil; ⁵Anton Pannekoek
Institute for Astronomy, University of Amsterdam, the Netherlands
email: mshultz@eso.org

Abstract. We present our analysis of 6 years of ESPaDOnS spectropolarimetry of the magnetic β Cep star ξ^1 CMa (B1 III). This high-precision magnetometry is consistent with a rotational period $P_{\text{rot}} > 40$ yr. Absorption line profiles can be reproduced with a non-rotating model. We constrain R_* , L_* , and the stellar age via a Baade-Wesselink analysis. Spindown due to angular momentum loss via the magnetosphere predicts an extremely long rotational period if the magnetic dipole $B_d > 6$ kG, a strength also inferred by the best-fit sinusoids to the longitudinal magnetic field measurements B_Z when phased with a 60-year P_{rot} .

Keywords. stars: circumstellar matter, stars: magnetic fields, stars: rotation

1. Magnetometry

The B0.7 IV β Cep star ξ^1 CMa was detected as magnetic by FORS1 (Hubrig et al. 2006) and later confirmed by ESPaDOnS (Silvester et al. 2009). Further FORS1/2 observations yielded a 2.2 d period (Hubrig et al. 2011), however many stars for which periods were reported based on FORS1/2 data were not confirmed to be magnetic by ESPaDOnS data (Shultz et al. 2012), casting doubt on the reliability of period analysis using FORS1/2 data. An analysis of an earlier, smaller ESPaDOnS data-set was presented by Fourtune-Ravard et al. (2011), who found $P_{\text{rot}} \sim 4$ d.

The ESPaDOnS dataset has grown to 34 Stokes V spectra obtained from 2008 to 2014, with annual clusters of up to 16 spectra separated by hours to days, and uniform integration times of 240 s. Sharp spectral lines and high SNR (>440) lead to a mean error bar in B_Z of $< \sigma_B > = 6$ G in the single-spectrum least-squares deconvolution (LSD) profiles (Kochukhov et al. 2010). B_Z declines smoothly from 340 ± 17 G in 2009 to 251 ± 3 G in 2014 (see Fig. 1). There is a systematic error (~ 15 G) due to a small variation of B_Z with the pulsation period. We have also located 2 archival MuSiCoS observations obtained in 2000. In both spectra, the polarity of the Zeeman signature is negative, with $B_Z = -137 \pm 32$ G. *The minimum P_{rot} compatible with the combined datasets is 40 years: the longest rotation period yet found for an early B-type star.*

2. Analysis

We used the Si III 455.3 nm line to derive $v \sin i$ and macroturbulence ζ , performing a goodness of fit test using a grid of synthetic line profiles including radial pulsation (Saesen et al. 2006), finding $v \sin i < 6$ km/s, and $\zeta = 20 \pm 3$ km/s, i.e. the line profile can be reproduced with a non-rotating model.

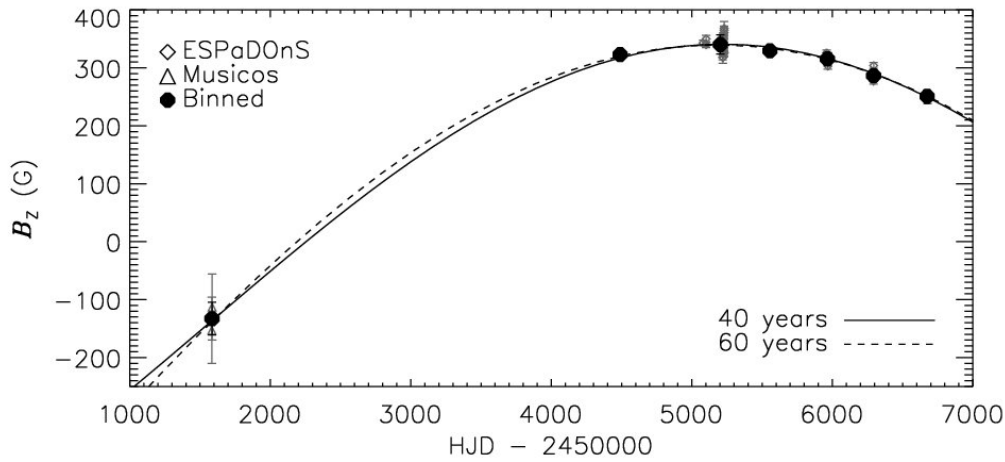


Figure 1. Longitudinal magnetic field B_z measurements as a function of time. Measurements binned by epoch are in solid black circles.

To constrain the stellar parameters we performed a Baade-Wesselink analysis, combining the pulsation phase-dependant variations in integrated light (via Hipparcos photometry), radial velocity (via ESPaDOnS and CORALIE spectra, Saesen et al. 2006), and T_{eff} (via EW ratios for various elemental ionic species). $\langle T_{\text{eff}} \rangle = 25.9 \pm 0.1$ kK, with a ± 0.5 kK amplitude variation. With the bolometric correction of Nieva (2013), the model light curve reproduces the photometric variability with $R = 8.7 \pm 0.7 R_{\odot}$, leading to $\log L/L_{\odot} = 4.46 \pm 0.07$, where the uncertainty is largely a function of the error in the Hipparcos parallax. Isochrones (Ekström et al. 2012) indicate an age of 12.6 Myr.

Magnetic stars shed angular momentum via the torque applied to the stellar surface by the corotating magnetosphere (Weber & Davis 1967; ud-Doula et al. 2009). Can a magnetic early B-type star spin down to $P_{\text{rot}} > 40$ yrs after only 12.6 Myr? We computed the Alfvén radius (ud-Doula & Owocki 2002) based on the mass-loss rate and wind terminal velocity predicted by the recipe of Vink et al. (2001) using the stellar parameters above, for values of B_d between 1.2 and 8.0 kG, and from this the spindown time (ud-Doula et al. 2009), assuming initially critical rotation. If $B_d > 6$ kG, $P_{\text{rot}} > 40$ yrs can be achieved; if $P_{\text{rot}} = 60$ yrs, the best-fit sinusoid to B_z implies $B_d \sim 6$ kG.

References

- Ekström, S., Georgy, C., Eggenberger, P., et al. 2012, *A&A* 537, A146
- Fourtune-Ravard, C., Wade, G. A., Marcolino, W. L. F., et al. 2011, in C. Neiner, G. Wade, G. Meynet, & G. Peters (eds.), *IAU Symposium*, Vol. 272 of *IAU Symposium*, pp 180–181
- Hubrig, S., Briquet, M., Schöller, M., et al. 2006, *MNRAS* 369, L61
- Hubrig, S., Ilyin, I., Schöller, M., et al. 2011, *ApJ* 726, L5
- Kochukhov, O., Makaganiuk, V., & Piskunov, N. 2010, *A&A* 524, A5
- Nieva, M.-F. 2013, *A&A* 550, A26
- Saesen, S., Briquet, M., & Aerts, C. 2006, *Communications in Asteroseismology* 147, 109
- Shultz, M., Wade, G. A., Grunhut, J., et al. 2012, *ApJ* 750, 2
- Silvester, J., Neiner, C., Henrichs, H. F., et al. 2009, *MNRAS* 398, 1505
- ud-Doula, A. & Owocki, S. P. 2002, *ApJ* 576, 413
- ud-Doula, A., Owocki, S. P., & Townsend, R. H. D. 2009, *MNRAS* 392, 1022
- Vink, J. S., de Koter, A., & Lamers, H. J. G. L. M. 2001, *A&A* 369, 574
- Weber, E. J. & Davis, Jr., L. 1967, *ApJ* 148, 217